GW+GRB+KN Q&A

Gravitational wave

LIGO detection of GW170817 Binary neutron star inspiral Time 12:41:04.4 UTC

Primary mass m_1 $1.36-1.60 M_{\odot}$ $1.17 - 1.36 M_{\odot}$ Secondary mass m_2 $1.188^{+0.004}_{-0.002} M_{\odot}$ Chirp mass \mathcal{M} Mass ratio m_2/m_1 $2.74^{+0.04}_{-0.01}M_{\odot}$ Total mass $m_{\rm tot}$ Radiated energy $E_{\rm rad}$ $> 0.025 M_{\odot} c^2$ 40^{+8}_{-14} Mpc Luminosity distance $D_{\rm L}$ Viewing angle Θ Using NGC 4993 location

$$\mathcal{M} = (m_1 m_2)^{3/5} (m_1 + m_2)^{-1/5}$$

0.7 - 1.0

 $\leq 55^{\circ}$

 $\leq 28^{\circ}$



Gravitational wave

- Low signal of *Virgo* provided strong constraint on sky position.
- Localized to within
 28 sq. deg. (>600 sq. deg 0° in previous GW detections)





LIGO + Virgo (2017)

Gravitational wave

- Distance of 44+3-7 Mpc was determined with only GW from "standard siren," independent from current distance techniques (e.g., standard candles).
- Together with redshift from EM localization (GW does not provide redshift), Hubble constant is 70+12-8 km/s/Mpc.



Freedman (2017)

Gamma-ray

- 1.7 seconds later, short GRB, GRB 170817A independently detected by *Fermi* and *INTEGRAL*.
- Unusually faint compared to other short-duration GRBs.
- Short GRBs long suspected to result from NS-NS mergers.



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X-ray + radio

- If the system launches relativistic jets as the detection of short GRB suggests, synchrotron radiation mainly in X-ray and radio is expected.
- Delayed (~10 days) and rising X-ray emission from *Chandra* interpreted as a jet with an opening angle ~15 ° and ~28° off axis.



Optical + NIR

- After GW+GRB, race is on to localize the source. Swope at Las Campanas won the price.
- SSS17a/AT2017gfo
- RA of 13h in August, setting, ~3hr from Sun!
- Optical observations
 10.86 hr after merger
 provided localization in
 outskirts of NGC 4993.
- Tully-Fisher distance 41.1
 Mpc, redshift z=0.0097.



Coulter et al. (2017)

Optical + NIR

- Evolution follows theoretical "kilonova," powered by radioactive decay of r-process nuclei synthesized in neutron rich merger ejecta.
- Initial rapid peak in optical, followed by broad peak in NIR.



Cowperthwaite et al. (2017)

Optical + NIR

• Lines are broad (high velocity) and some light r-process elements were identified in the spectra.



Kilonova

- NS-NS or NS-BH mergers should produce neutron-rich radioactive species, whose decay should result in a faint transient, "kilonova," main source of stable r-process elements in the Universe.
- Ejecta rich in heavy r-process elements expected to peak in the NIR because of high optical opacity.



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Kilonova

- SSS17a is not the first observed: GRB130603B + kilonova
- z=0.356 or 1800 Mpc or distance modulus 41 mag
- "Kilo"nova is ~1000x
 brighter than nova or
 —15 absolute mag
- Magnitude of 26 mag, exceedingly difficult observation.



Kilonova



What's left?

- Inferred mass of the "red kilonova" large, suggests that it is produced in post-merger disk wind and the remnant collapsed rapidly => black hole.
- If remnant is a hot neutron star that survived longer than 10s ms, neutrino irradiation would reduce neutron fraction => "blue kilonova"



Kasen et al. (2017)



LIGO + Virgo (2017)

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Summary

- During merger, tidal force strips matter, form torus.
- Red: neutron-rich heavy r-process ejecta produces optical-NIR emission peaking in NIR.
- Blue: neutron-poor light r-process ejecta produces optical emission.
- Black: collimated jet with synchrotron radiation emitting radio and X-ray. Faint short GRB due to viewing angle.



We got lucky!

- Virgo was online.
- It is the most nearby short-duration GRB.
- "Favorable" viewing angle to obtain lots of information.
- It would have been behind the Sun in 1.5 months.

